

The potential of Ca II K observations for solar activity and variability studies

Ilaria Ermolli¹, Theodosios Chatzistergos^{1,2}, Natalie A. Krivova² and
Sami K. Solanki^{2,3}

¹INAF – Osservatorio Astronomico di Roma, Monte Porzio Catone, Italy

²Max-Planck-Institut für Sonnensystemforschung, Göttingen, Germany

³School of Space Research, Kyung Hee University, Republic of Korea

Abstract. Several observatories around the globe started regular full-disc imaging of the solar atmosphere in the Ca II K line in the early decades of the 20th century. These observations are continued today at a few sites with either old spectroheliographs or modern telescopes equipped with narrow-band filters. The Ca II K time series are unique in representing long-term variations of the Sun's chromospheric magnetic field. However, meaningful results from their analysis require accurate processing of the available data and robust merging of the information stored in different archives. This paper provides an overview of the historical and modern full-disc Ca II K observations, with focus on their quality and the main results obtained from their analysis over the last decade.

Keywords. Sun: activity, plages, Sun: chromosphere, Sun: magnetic fields

1. The Ca II K line

The visible spectrum of the Sun, the brightest star in our sky, contains millions of absorption lines. Two of the deepest and broadest such lines are the resonance doublet lines of singly-ionized calcium at 3933 Å and 3968 Å. First observed by Joseph von Fraunhofer in 1814, these lines were designated as K and H, respectively, in his catalogue of prominent absorption features. Linsky & Avrett (1970) consider that “The significance of these lines lies in the fact that they are the only resonance lines in the visible spectrum from the dominant stage of ionization in the upper photosphere and lower chromosphere of an abundant element. Consequently, these lines are very opaque with line centre optical depths of order 10^7 in the photosphere and 10^4 at the temperature minimum. The entire cores of the lines are thus chromospheric and can serve as convenient probes of the structure and physical properties of the chromosphere.”

Figure 1 shows two spatially-averaged intensity profiles of the Ca II K line measured in a quiet Sun and a plage region superimposed on each other. The line profile is characterized by the two peaks K_{2V} and K_{2R} towards the Violet or Red part of the spectrum relative to the line centre, respectively, the reversal at the line centre K_3 , and the two secondary minima K_{1V} and K_{1R} . These line features result from sampling of different parts in the inner solar atmosphere. From the line wings towards the line centre, the line samples the photosphere with increasing height until the temperature minimum, where the K_1 minima are formed. The decreasing temperature with height in the photosphere produces a regular absorption line. Closer to the line core between the K_1 minima, the line samples the low chromosphere where the temperature increases with height, the source function essentially follows the Planck function and the line thus gets brighter with increasing height. This is the case nearly up to the atmospheric heights where the K_2 peaks are formed, where the source function becomes smaller than the Planck function

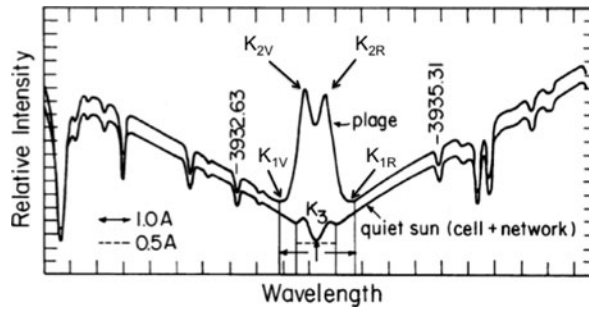


Figure 1. Ca II K line profile for quiet Sun (lower curve) and plage region (upper curve). Adapted from Skumanich *et al.* (1984).

and absorption dominates again, causing the K_3 minimum. These line features appear qualitatively the same for both the quiet Sun and the plage region, but for the latter the K_2 peaks and the K_3 minimum display a significant increase of the contrast relative to the quiet region. This shows that the chromospheric heating is greatly enhanced in regions where the magnetic field is stronger.

It is worth noting that all individual line features introduced above occur within a 1 \AA interval. Ermolli *et al.* 2010 studied the intensity response function for Ca II K observations taken with different filters and showed that for data taken with a 1 \AA bandwidth filter centred at the K_3 , 58–82% of the contribution at disc centre is from atmospheric heights below 500 km. For a 2.5 \AA wide filter, i.e. similar to that of the present-day HINODE SOT/BFI (Tsuneta *et al.* 2008) and broad Ca II K Rome-PSPT (Ermolli *et al.* 2007) observations, this contribution is as high as 84–94%. Besides, for observations with a 1 \AA bandwidth filter centred in the red wing of the Ca II K line, more than 99.6% of the contribution is from atmospheric heights below 160 km.

2. Overview of historical and modern observations

In the early 1890s, Henri Alexandre Deslandres and George Ellery Hale independently put into practice an instrument to register full-disc photographic images of the chromosphere at the Ca II K and other spectral lines.

This new instrument, called spectroheliograph, comprises two main parts. The first part is a system consisting of either one or two mirrors directing the Sun's light to an objective lens that focuses the image of the Sun at the entrance window of the second part of the instrument. This includes a prism or a diffraction grating and a monochromator to allow the passage of a single wavelength through the exit window of the instrument. Both the entrance and exit windows are slit shaped. After the exit window the photographic material used to record the solar image, most commonly being a photographic plate, or film, and more recently CCD cameras, is placed. A proper displacement of the objective lens allows the full-disc image of the Sun to pass in front of the instrument's entrance slit. With the motion of the Sun's image at the entrance slit, a proper motion of the photographic material (or of the recording CCD) allows a full-disc image of the Sun to be recorded at the wavelength the instrument is operated.

Following the first two models at the Paris and Kenwood observatories, several spectroheliographs were built and put in regular operation at different sites, e.g. in 1904 at the Kodaikanal, in 1915 at the Mt Wilson, in 1917 at the Mitaka, in 1926 at the Coimbra, and in 1931 at the Arcetri observatories. The spectroheliographs at the Coimbra and Paris observatories, each equipped with a 45 cm coelostat and 25 cm objective lens, are

still operated to this day. In 2008 and 2002, respectively, the photographic plates were replaced with CCD cameras.

In the 1960s, regular full-disc Ca II K observations started at several sites also with telescopes equipped with narrow band Lyot-type and interference filters, e.g. in 1964 at the Rome, 1968 at the Kandilli, and in 1971 at the Big Bear observatories.

A large fraction of the historical Ca II K observations is still stored in its original form, i.e., in photographic archives, some of which have recently been digitized, e.g. the Arcetri (1931–1974), Kodaikanal (1904–2007), Kyoto (1926–1969), Mitaka (1917–1974), and Mt Wilson (1915–1985) ones.

At present, routine full-disc Ca II K observations are performed with modern instruments at the Baikal, Brussels, Chrotel-Teide, Kanzelhöhe, Mitaka, Pic du Midi, Rome, San Fernando, and Valašské Meziříčí Observatories. These series started in 2003, 2012, 2009, 2010, 2015, 2007, 1996, 1986, and 2013, respectively.

Further information on existing Ca II K archives can be found in e.g. Chatzistergos (2017) and Chatzistergos *et al.* (2018a).

3. Scientific programs

Since early observations, regular Ca II K data have been used to monitor the evolution of the most prominent disc features, in particular of plages and prominences. Besides, they have long served as diagnostics of the solar chromosphere. For example, Ca II K data have been used to understand the formation of the Ca II K line and to produce an atmosphere model of the broad height range sampled by the line. Indeed, the temperature structure of the semi-empirical atmosphere models by Vernazza *et al.* (1981) and Fontenla *et al.* (1993) were also tested on the available Ca II K observations. Over time, analysis of the differences between spatially-averaged profiles and high-resolution observations of the Ca II K line in different disc regions have also given hints to the strongly inhomogeneous and dynamic nature of the chromosphere. However, the most prominent aspect of the Ca II K line is its ability to provide information about solar magnetism, due to the large increase in the intensity of the K_3 line core when sampling bright magnetic regions (plage). The potential of the Ca II K observation to serve as proxy of the magnetic field, which was first noticed by Babcock & Babcock (1955), has inspired many scientific programs aimed at determining the evolution of the solar magnetic field. To this aim, the position, dimensions, number, type, and evolution of the different magnetic regions observed on the solar disc were derived from the observations collected at several sites. The measurements of these various parameters were tabulated, graphically represented, and published in volume series by several observatories. Copies of the photographic observations were also produced and distributed to various research centres. Finally, the connection between the Ca II K emission and magnetism has long served to study the magnetic activity of stars other than the Sun.

The photometric observations obtained over the last two decades have also been used to analyse the radiative properties of the different disc features in time. For example, Ermolli *et al.* (2007) studied the photometric properties of plages over the activity cycle, from analysis of Rome-PSPT data obtained from 1998 to 2005, showing that selection effects associated with identification methods can produce significant differences in the results. They also showed that the plage contrast is a function of the selection method and the heliocentric angle, but also of the feature size, the activity level, and the content of the analysed images. Besides, Ermolli *et al.* (2010) analysed the radiative emission of various types of solar features, such as quiet Sun, enhanced network, plage, and bright plage regions, studying the dependence of the measured radiative emission on the

filter bandpass and comparing results from observations with those derived from spectral synthesis performed on semi-empirical atmosphere models in the literature.

Over the last decade, digitization of some prominent photographic archives has initiated extensive exploitation of the Ca II K series for studies of the long-term variation of the chromospheric magnetic field and a variety of retrospective analyses. To date, historical Ca II K data have been mainly used to produce e.g. time series of the disc coverage of plage and network regions, i.e. of the other obvious magnetic features on the solar disc besides sunspots. For example, Ermolli *et al.* (2009b), Foukal *et al.* (2009), Tlatov *et al.* (2009), and Chatterjee, Banerjee & Ravindra (2016) presented plage area time series derived from analysis of different archives with distinct image processing methods. The annual mean time series presented by these authors show some similarities between independent analyses, but also display significant differences in the overall long-term trends and rank ordering of activity cycle amplitudes. Besides, since Ca II K observations track the spatial distribution of magnetic flux distribution over the disc, Harvey (1992), Ermolli *et al.* (2009a), Chatterjee, Banerjee, Ravindra (2016) produced butterfly diagrams of plage regions identified on the Ca II K observations of the Mt Wilson, Arcetri, and Kodaikanal observatories, respectively. These diagrams complement similar data obtained from sunspot observations and extend butterfly diagrams derived from magnetograms to times when such data were not available. Moreover, to study the magnetic field at different heliographic latitudes and the toroidal and poloidal components of the solar cycle, Sheeley, Cooper & Anderson (2011) and Pevtsov *et al.* (2016) produced Carrington maps from Ca II K observations of the Mt Wilson and Kodaikanal archives.

Finally, an analysis of the supergranulation scale variation based on present-day Rome-PSPPT Observations and historical Kodaikanal data was presented by e.g. Ermolli *et al.* (2003) and Chatterjee *et al.* (2017), respectively. Network area measurements were presented by Foukal & Milano (2001) and Priyal *et al.* (2014).

Understanding the variation of plage and network properties from analysis of Ca II K observations has a broader impact than just describing two of the most prominent solar disc features. Indeed, such measurements may be useful for constraining models of magnetic flux transport on the surface of the Sun during multiple sunspot cycles, and for estimating the long-term contribution of facular emission to the total and spectral solar irradiance. Accurate and significant analysis of long-term Ca II K series has thus a lot of potential for new knowledge in solar physics.

4. Comparison among series

All studies listed above confirm some well-known characteristics of past cycles, e.g. they all report that solar cycle 19 showed remarkably high plage coverage and broad latitudinal distribution of active regions. They also agree on the overall increase of the solar activity over the first half of the 20th century and decrease over the last decades. However, comparison of results presented in the literature also clearly shows that most of the reported findings are inherently inconsistent, due to large differences in absolute values and details in the detected short- and long-term trends. These differences affect results derived from distinct time series, as well as findings from the same archives obtained in different studies.

As an example, in Fig. 2 we compare the power spectrum of three published plage area time series using the Morlet wavelet. In particular, we consider plage area values derived by Foukal (1996), Tlatov *et al.* (2009), and Chatterjee *et al.* (2016) from the Mt Wilson and Kodaikanal archives and compare them with results for the sunspot number series (SN); focus is on the periods between 2 to 32 years. The dominant feature in all panels

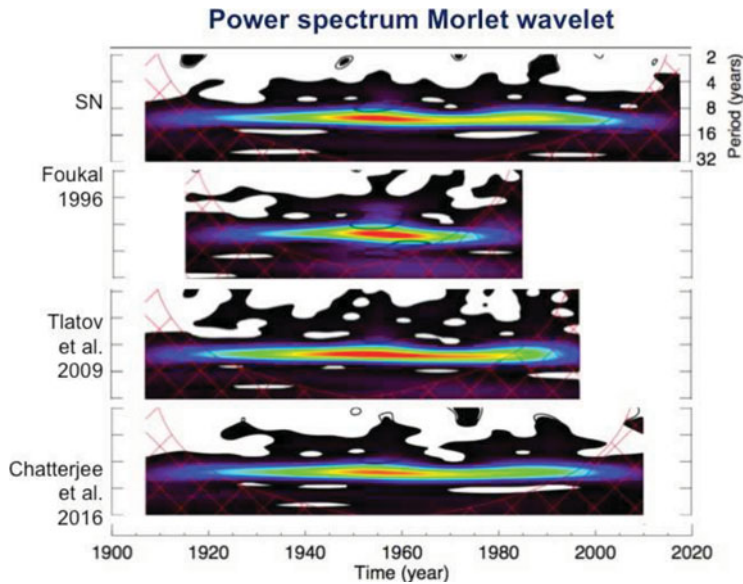


Figure 2. Example of the power spectrum of plage area timeseries reported in the literature using the Morlet wavelet. From top to bottom, power spectrum of the sunspot number series (SN) and of the plage area values by Foukal (1996), Tlatov *et al.* (2009), and Chatterjee *et al.* (2016), from analysis of the Mt Wilson and Kodaikanal series. For the sake of clarity the panels only show the power spectrum for periods between 2 to 32 years. Black to red show low to high values, respectively.

is obviously the 11 year cycle, and cycle 19 is the strongest one in all records. However, some differences among the series in the modulation of the high- and low-frequency variations are also seen. It is worth noting that the considered plage area values were independently derived from analysis of different datasets, but the studies by Tlatov *et al.* (2009) and Chatterjee *et al.* (2016) are based on images from different digitization of the same archive.

The discrepancy in published results can be mostly attributed to differences in observation set-ups, to different samples of considered data, and to the data processing applied by the various authors. As shown by Chatzistergos *et al.* (2018b), simple visual inspection reveals considerable variance between the images taken on the same days at different sites, because of instrumental and operational characteristics. Furthermore, Ermolli *et al.* (2009b) showed that the image quality and contents of three prominent Ca II K archives varied with time, probably due to instrument changes, aging of the spectroheliographs used, degradations of the photographic material, and changes in the observing programs. However, another important reason for the discrepancies in published results is the processing applied to the data. Historical observations suffer stronger geometrical distortions and photometric uncertainties than similar present-day data and hence demand a significantly more careful processing than modern observations to derive meaningful results from their analysis. It is worth noting that most of the results presented in the literature lack error estimates. Chatzistergos *et al.* (2018a) showed that with proper methods it is possible to derive consistent and accurate results e.g. of the plage coverage for images from different Ca II K archives including both historical and modern observations. See also the contribution by Chatzistergos *et al.* (2018b) in this volume.

5. Conclusions

A large quantity of full-disc Ca II K observations has been archived at different sites since the beginning of the 20th century. A few series are still being updated and extended using either old spectroheliographs or modern telescopes equipped with narrow-band filters. Some historical observations have been digitized and made available to the community. These time series are particularly important because they carry information on the chromospheric magnetic field over a time span exceeding 110 years. Indeed, apart from sunspots, chromospheric plage and network are the next obvious magnetic features on the solar disc and they are clearly spotted in the Ca II K observations.

Analyses of Ca II K time series show both cyclic and long-term variations of plage and network properties. Results presented in the literature from analysis of available data are reasonably in agreement, since they all show some known features of past activity cycles and of long-term solar variation, but also inherently inconsistent, due to differences in the employed data and the methods used to process the observations.

Understanding variations of both plage and network properties has a broader impact than just describing two of the most prominent solar surface features. However, meaningful results on the long-term variation of the magnetic field and on solar variability from analysis of Ca II K datasets require accurate processing of the available observations and merging of the information stored in different archives. In fact, analysis of multiple archives is critical for detecting and correcting instrumental trends, both systematic and otherwise. Besides, study of present-day observations is also fundamental for optimizing and evaluating the accuracy of the processing applied to historical datasets.

References

- Babcock, H. W., & Babcock, H. D. 1955, *ApJ*, 121, 349
- Chatterjee, S., Banerjee, D., & Ravindra, B. 2016. *ApJ*, 827, id.87
- Chatterjee, S., Mandal, S., & Banerjee, D. 2017, *ApJ*, 841, id.70
- Chatzistergos, T. 2017, Analysis of historical solar observations and long-term changes in solar irradiance (Uni-edition), Göttingen
- Chatzistergos, T., Ermolli, I., Solanki, S. K., & Krivova, N. A. 2018a, *A&A*, 609, A.92
- Chatzistergos, T., Ermolli, I., Solanki, S. K., & Krivova, N. A. 2018b, *Proc. IAU Symposium* 340
- Ermolli, I., Berrilli, F., & Florio, A. 2003, *A&A*, 523, 857
- Ermolli, I., Criscuoli, S., Centrone, M., Giorgi, F., & Penza, V. 2007, *A&A*, 523, 305
- Ermolli, I., Marchei, E., Centrone, M., *et al.* 2009a, *A&A*, 499, 627
- Ermolli, I., Solanki, S. K., & Tlatov, A. G., *et al.* 2009b, *ApJ*, 698, 1000
- Ermolli, I., Criscuoli, S., Uitenbroek, H., *et al.* 2010, *A&A*, 523, A.55
- Fontenla, J. M., Avrett, E. H., & Loeser, R. 1993, *ApJ*, 406, 319
- Foukal, P. 1996, *GeoRL*, 23, 2169
- Foukal, P. & Milano, L. 2001, *GeoRL*, 28, 883
- Foukal, P., Bertello, L., Livingston, W. C., *et al.* 2009 *SolPhys*, 255, 229
- Harvey, K. 1992, *The solar cycle*, ASP Conference Series, 27, 335
- Linsky, J. L. & Avrett, E. H. 1970, *PASP*, 82, 169
- Pevtsov, A. A., Virtanen, I., Mursula, K., Tlatov, A., & Bertello, L. 2016, *A&A*, 585, A.40
- Priyal, M., Singh, J., Ravindra, B., Priya, T. G., Amareswari, K. 2014 *SolPhys*, 289, 137
- Sheeley, N. R., Jr., Cooper, T. J., & Anderson, J. R. L. 2011, *ApJ*, 730, 51
- Skumanich, A., Lean, J. L., Livingston, W. C., & White, O. R. 1984, *ApJ*, 282, 776
- Tlatov, A. G., Pevtsov, A. A. & Singh, J. 2009 *SolPhys*, 255, 239
- Tsuneta, S., Ichimoto, K., Katsukawa, Y., *et al.* 2009 *SolPhys*, 249, 167
- Vernazza, J. E., Avrett, E. H., & Loeser, R. 1981, *ApJS*, 45, 635